

Characterizing CO Fourth Positive Emission in Young Circumstellar Disks

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ABSTRACT

Carbon Monoxide is a commonly used IR/sub-mm tracer of gas in protoplanetary disks. We present an analysis of ultraviolet CO emission in *HST*-COS spectra for 12 Classical T Tauri stars. Several ro-vibrational bands of the CO $A^1\Pi - X^1\Sigma^+$ (Fourth Positive) electronic transition system are spectrally resolved from emission of other atoms and H₂. The CO $A^1\Pi v'=14$ state is populated by absorption of Ly α photons, created at the accretion column on the stellar surface. For targets with strong CO emission, we model the Ly α radiation field as an input for a simple fluorescence model to estimate CO rotational excitation temperatures and column densities. Typical column densities range from $N_{CO} = 10^{18} - 10^{19}$ cm⁻². Our measured excitation temperatures are mostly below $T_{CO} = 600$ K, cooler than typical M-band CO emission. These temperatures and the emission line widths imply that the UV emission originates in a different population of CO than that which is IR-emitting. We also find a significant correlation between CO emission and the disk accretion rate \dot{M}_{acc} and age. Our analysis shows that ultraviolet CO emission can be a useful diagnostic of CTTS disk gas.

Subject headings: protoplanetary disks — stars: individual (RECX 15, DM Tau, DE Tau, HD 135344B, DR Tau, DK Tau, DF Tau, LkCa 15, GM Aur, HN Tau, RECX 11, V4046 Sgr)

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1. Introduction

Classical T Tauri stars (CTTSs) provide the opportunity to witness the birth of extra-solar planetary systems. Their disks of gas and dust are the environments in which Jovian planets form (Dullemond & Monnier 2010), a process that we are only beginning to understand. From the ground, carbon monoxide (CO) is the easiest molecule to observe in emission from these disks, and has been studied from near-IR to mm wavelengths (Salyk et al. 2008; Dutrey et al. 2008; Sturm et al. 2010). Such observations sample regions from $r_{in} < 1$ AU (Najita et al. 2003) to $r_{out} > 800$ AU (Pi  tu et al. 2007). The second most abundant molecule in the galaxy (after H₂), CO plays a significant role in the formation of planetary systems. As H₂ does not emit strongly in the IR, CO is an important tracer of disk gas. CO abundance varies greatly in different environments throughout the Milky Way. In the diffuse ISM, the CO/H₂ ratio ranges from 10^{-7} to 10^{-6} (Burgh et al. 2007); in dense clouds, it can be 10^{-4} (Lacy et al. 1994), while comets in the solar system exhibit a ratio of ~ 10 (Lupu et al. 2007; Feldman et al. 2009). During the process of star and planet formation, the CO and H₂ may be selectively removed from or added to the gas phase in a way that changes the CO/H₂ ratio (Vidal-Madjar et al. 1994; France et al. 2011a). By probing CO in protoplanetary disks, we can constrain models of planetary system formation.

Electronic transitions of CO have recently been identified in the *Hubble Space Telescope* Cosmic Origins Spectrograph (*HST*-COS) spectra of several CTTSSs (France et al. 2011a). Several vibrational bands of the $A^1\Pi - X^1\Sigma^+$ (Fourth Positive) system were measured in emission for V4046 Sgr and RECX 11, and in absorption in HN Tau. Lyman α (Ly α) photons produced in the accretion funnel flow excite the CO, populating its $A^1\Pi v' = 14$ state; several of the resulting (14 - 2) through (14 - 12) transitions are detected between 1280 Å and 1720 Å. Emission from the (0 - 1) 1600 Å and (0 - 3) 1713 Å bands also appears, potentially due to C IV photoexcitation of CO. The shape and strength of these emission bands vary with temperature (T_{CO}) and column density (N_{CO}) such that these parameters can be determined. Circumstellar CO can be observed in absorption as well; a collisionally-excited distribution of gas at $T_{CO} = 500 \pm 200$ K is able to reproduce several of the ($v' - 0$) absorption bands in the HN Tau spectrum for $\log_{10}(N_{CO}) = 17.3 \pm 0.3$ cm $^{-2}$. France et al. (2011, in press) also find CO in absorption in AA Tau, with $\log_{10}(N_{CO}) = 17.5 \pm 0.5$ cm $^{-2}$.

We extend this work by conducting a survey of COS spectra for 12 CTTSs in greater detail. Our aim in this paper is to characterize T_{CO} and N_{CO} in a variety of CTTSs, so as to understand correlations between the UV-emitting CO (UV-CO) and the disk properties that characterize these systems. The UV-CO can also be compared with CO previously detected in these disks at IR and mm wavelengths. We model the Ly α radiation field assuming absorption from an outflow and interstellar H I, using this as an input for simple

fluorescence models. These models enable us to estimate T_{CO} and N_{CO} for targets with high signal-to-noise (S/N) CO emission. In §2 we detail the COS observations, then discuss the CO emission in section §3. In §4 we summarize our model for the Ly α profiles as well as the CO fluorescence model. In §4.2 we discuss the results of this analysis, and our ability to constrain the temperatures and column densities of CO in each of the targets. We also compare these empirically derived properties with other previously measured parameters from the literature.

2. Observations

We investigated a sample of 12 CTTSs observed by *HST*-COS; the observational parameters are listed in Table 1 in order of observation date. These were the first CTTSs observed with COS. Spectra from several of these targets have already been presented in the literature. Ingleby et al. (2011) studied accretion rates in RECX 11 and RECX 1, and Yang et al. (2011) demonstrated Ly α photon absorption and re-emission by H₂ in V4046 Sgr and DF Tau. France et al. (2011b) characterized the continuum emission of V4046 Sgr and DF Tau, while France et al. (2011a) identified CO Fourth Positive emission and absorption in the spectra of V4046 Sgr, RECX 11, and HN Tau. This paper builds upon the work of France et al. (2011a), advancing the models for analyzing the CO emission of a survey of CTTSs.

Each spectrum consists of 2 - 4 orbits with the G130M and G160M gratings (Osterman et al. 2011), yielding a moderate resolving power of $\lambda/\delta\lambda \sim 18,000$. As there is no field stop in the instrument, the entire objects were sampled without appreciable spatial resolution relative to the size of the disks. We used CALCOS, the COS calibration pipeline, to reduce the data and then created coadded spectra from 1140 Å to 1760 Å. The data reduction pipeline creates inaccurately small error bars in low count regions of the spectrum (Froning et al. 2011). This gives excessive weight to continuum points in the least squares fits of the emission feature. We empirically determined that this was generally the case for data points with errors more than three times smaller than the mean 1σ continuum error. Any errors smaller than this limit were thus set to the limit value.

Most targets belong to the η Chamaeleontis and Taurus-Auriga star formation regions, with properties described in Appendix A. Table 2 lists a variety of target parameters relevant to this work, including visual extinction (A_V), accretion rate (\dot{M}_{acc}), inclination, and age.

3. Ly α -excited CO Emission

CO UV emission appears in 7 of the 12 targets in our survey. We find strong CO (14 - v'') emission in four targets (V4046 Sgr, DF Tau, RECX 15, and DM Tau), and weaker but identifiable emission in RECX 11, HD135344 B, and DE Tau. The CO emission bands seen are listed in Table 3. The (14 - 1), (14 - 6), (14 - 9), and (14 - 11) bands are unidentifiable due to either low branching ratios or overlying emission from other molecules and ions. We list the branching ratios $B(v' - v'')$ for each band, calculated by summing the A values of each J transition in a given band, and dividing by the sum of all ($v' - v''$) transitions for a given v' . Despite the relatively low A values of the overlying (14 - 0) transitions (1.32×10^4 s $^{-1}$ for $J'' = 0$), there is sufficient Ly α flux to yield strong CO emission. C IV-pumped emission bands are also seen in several targets; we present this evidence in §5.3. The Ly α -excited emission is the primary subject of our analysis.

Figures 1 and 2 show the observed Ly α profiles and (14 - 3) CO bands for the sample. In this paper we use femto-erg flux units (FEFU, where $1 \text{ FEFU} = 1 \times 10^{-15} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ \AA}^{-1}$) as a standard flux unit. The geocoronal Ly α emission in the middle of the Ly α profile is set to zero to avoid confusion with target Ly α emission. For future discussion, we refer to the Ly α emission on the longer wavelength side of the interstellar absorption trough as the “red wing”, and the emission on the shorter wavelength side the “blue wing”. The CO plots in these figures also separate the CO bands into a “high- J ” and “low- J ” side, identified by the color of the continuum. The blueward CO emission comes from the lower- J transitions (0 - 15), which are closely spaced relative to the instrument resolution, forming a “bandhead”. The redward CO emission is composed of higher J transitions, which are well separated in several targets. There is a general correlation between the relative amount of flux in the Ly α blue and red wing and the relative amount of flux seen in the low- J and high- J CO emission. For example, in RECX 15 and DF Tau, weak blue wing Ly α emission correlates with the lack of low- J emission in their CO bands, whereas in V4046 Sgr we observe strong blue and red emission in both the Ly α and CO profiles. This will motivate our modeling, as discussed in §4.

The (14 - 3) band is mostly free of emission from other molecules or atoms; only the weak H₂ (4 - 2) P(19) line at 1316.55 Å appears in all targets. There is an H₂ (3 - 2) P(18) line at 1320.22 Å, but it appears strongly only in DF Tau. Overall, the low- J bandhead and separated high- J lines of the (14 - 3) band are uncontaminated, which is crucial for determining N_{CO} and T_{CO} . For this reason, we primarily use the (14 - 3) band for analysis in all targets. By branching ratios, the (14 - 4) band is stronger, however strong OI] (1355.71 Å) emission lines appear over the high- J portion of the band in most targets. Figure 3 shows the (14 - 4) band for our sample, with similar continuum denotations as the (14 - 3) band.

Only V4046 Sgr, RECX 15, and DM Tau are free from contamination on the high- J side. For these three targets we analyze both the (14 - 3) and (14 - 4) bands.

In V4046 Sgr, DF Tau, DM Tau, and RECX 15, the high- J lines are well separated. The high- J lines of DF Tau, RECX 15, and DM Tau appear at the width of the COS line-spread-function. In RECX 11 and HD 135344B the CO high- J lines are not distinguishable from one another, although whether this is physical (due to velocity broadening of close-in CO) or merely due to the lower S/N in the band is not clear. The Ly α profile of DE Tau is highly absorbed, yet the (14 - 3) features are evident to the eye. HN Tau, DR Tau, DK Tau, LkCa 15, and GM Aur show little CO emission above the continuum. Of those, GM Aur is the only target to have a Ly α profile that is mostly unabsorbed by circumstellar and interstellar H I.

4. Method and Modeling Results

4.1. Ly α and CO Fluorescence Model

France et al. (2011a) utilized a simple fluorescence model to calculate the temperature T_{CO} and column density N_{CO} for two CTTSs. However, they used the observed Ly α emission to excite the CO. In our full sample, the targets demonstrate the need for a more detailed estimation of the incident Ly α radiation; using the observed Ly α profiles as the incident radiation field cannot reproduce the observed CO band shape. For example, with DF Tau, more red wing Ly α emission than is seen is necessary to yield the strong CO (14 - 3) high- J line emission observed, relative to its weak low- J emission. We therefore introduce a straightforward model of the physical situation, shown in Figure 4, which we use to re-construct the incident radiation. Ly α emission is created at the accretion funnel flow near the proto-stellar surface, as well as in the photosphere itself. The plot in Figure 4 shows such a model Ly α profile for V4046 Sgr in red. Along our line of sight the Ly α is then absorbed by H I in an outflow (magenta), followed by interstellar gas (green). Our method consists of two steps: first, we create a series of model Ly α profiles that fit the observed profile after absorption by outflow and ISM H I. Second, we use each outflow-absorbed model Ly α emission in a CO fluorescence model to match the observed CO emission.

The observed Ly α emission is modeled as:

$$I_{obs} = C_{Ly\alpha} + I_{Ly\alpha} e^{-(\tau_{out} + \tau_{ISM})} \quad (1)$$

where $C_{Ly\alpha}$ is the continuum intensity near Ly α and the source Ly α profile $I_{Ly\alpha}$ is approx-

imated as the sum of two Gaussian emission components, parameterized by peaks I_1 and I_2 , as well as widths σ_1 and σ_2 . The first Gaussian is at rest relative to the star, while the second Gaussian component is centered at a velocity v_2 . The outflow absorber τ_{out} is characterized by a Voigt profile with N_{out} at a velocity v_{out} , and is necessary to match blue-shifted absorption seen in most of the targets. The ISM absorber τ_{ISM} is determined by a Voigt profile with N_{out} at the local standard of rest (LSR). A_V values from Table 2 and $R_V = 4.0$ were used to de-redden the data, according to the extinction curve from Cardelli et al. (1989). Seven parameters determine the two emission (stellar and accretion shock) and two absorption (interstellar and outflow) components. As a result, many parameter combinations exist that in the end fit the observed Ly α profile. However, only some will adequately produce the observed CO emission, once run through the fluorescence model.

Our CO fluorescence model uses a method similar to that implemented by France et al. (2011a). The ground rotational levels are populated thermally with T_{CO} , using Dunham coefficients from George et al. (1994) to calculate ground state energy values. The Ly α flux absorbed by the (14 - 0) band is re-emitted according to the branching ratios for the (14 - v'') transitions. However, instead of using the observed Ly α flux, we excite the CO with the outflow-absorbed profile. For each Ly α profile produced in the first step, we produce a model CO emission band to fit to the data. The variable parameters in the CO model are the gas rotational temperature (T_{CO}), column density (N_{CO}), and the velocity broadening of the emitting gas (Δv_{broad}). This Δv_{broad} value is the full width at half maximum (FWHM) of the Gaussian emission line, assuming the broadening is primarily due to orbital motion. For a continuum, we use the average continuum level from nearby areas of the spectrum. Assuming that the electronic ground state population can be described by a thermalized population, the resulting CO emission is primarily dependent on N_{CO} and T_{CO} . As in France et al. (2011a), we make no assumptions about the spatial distribution of the CO around the star. Due to the low transition probabilities of the (14 - 0) band ($\sim 10^4$ s $^{-1}$), optical depth effects are minimal; therefore, N_{CO} and the filling fraction of the CO are degenerate.

There are limitations to our assumption of a single CO rotational temperature, as multiple temperatures are possible throughout the disk. The gas disk temperature varies based on height in the disk and distance from the star. As the entire disk is sampled in the COS aperture, CO may be observed at multiple temperatures. However, a single temperature model does enable us to determine general characteristics of the CO in a large sample with minimal parameters. Future models of individual targets should take into account the effects of density and temperature on the radiative transfer of the Ly α radiation and fluorescent emission.

4.2. Results

In half of the positive detection targets, the uncertainties on N_{CO} are less than 0.5 dex and the uncertainties on T_{CO} are less than ~ 200 K. High S/N in both the Ly α and CO profile are required for this. In most of the positive detections, values for $\log_{10}(N_{CO})$ are between 18 and 19 and temperatures lie between 100 and 600 K.

4.2.1. Ly α Model Fits

The top panel of Figure 5 shows representative Ly α model results for DF Tau. Several model Ly α profiles are plotted over the observed Ly α profile. The red, blue, and green lines correspond to the maximum, average, and minimum Ly α emission profiles after being absorbed by the outflow H I of each respective model. These are displayed to convey the variety of incident radiation profiles that, when used in the CO fluorescence model, produce CO emission that matches the data. The magenta line shows a typical outflow + ISM-absorbed profile that fits the Ly α profile. In the case of DF Tau the entire blue wing is absorbed, which we attribute to a non-gaussian absorber velocity profile; therefore we only fit the red wing. The blue wing is generally fitted between 1210 Å and 1214 Å, while the red wing is fitted from 1217 Å to 1220 Å. We did not use two Ly α emission components for DF Tau and RECX 11, as a second component did not significantly improve the fit. Strong H₂ absorption exists in the Ly α wings of V4046 Sgr, DF Tau, RECX 15, and DM Tau (Yang et al. 2011); a spline interpolation of the profile was used to minimize their influence on the fits. The H₂ absorption may be present in the Ly α profile incident on the fluoresced CO, but the absorption has a negligible effect on the resulting CO fluorescence.

We searched a grid of the parameters to find the best-fitting combinations. A “good” fit was determined by minimizing χ^2 . Fits were excluded with χ^2 greater than 95.4 % (2σ) from the peak of the χ^2 probability distribution. In Tables 4 and 5 we list parameter values of the best-fitting Ly α profiles. Since our model Ly α profiles must not only fit the observed Ly α emission but also produce the appropriate CO emission through our fluorescence model, we only include models that create CO emission that fits the (14 - 3) band. The Ly α parameters listed are determined by taking the average and standard deviation of each parameter for all models. The broad range of values for the emission components reflects the variety of emission profiles that in combination with different absorbing profiles match the data. Continuum values are determined by averaging emission-free ranges of the surrounding spectrum. The “outflow” absorber velocities for DM Tau are positive, possibly due to predominantly inflowing material. Overall, our combined values for N_{ISM} and N_{out} are similar to those measured previously in CTTSs (Walter et al. 2003; Herczeg et al. 2006).

4.2.2. CO Model Fits

The middle panel of Figure 5 shows a typical CO model fit to the (14 - 3) band, along with its value for N_{CO} and T_{CO} . However, this N_{CO} and T_{CO} combination is the best fit for only one of our model Ly α profiles. The bottom panel shows the distribution of best fit T_{CO} vs. N_{CO} for all the best fit Ly α profiles, to demonstrate the range of possibilities enabled by our model. Each point corresponds to a unique set of Ly α profile, N_{CO} , and T_{CO} that in combination adequately reproduce the observed Ly α and CO emission. The error bars for each fit are not included, as they are smaller than the overall distribution of points. Ly α and CO must be significantly detected in order to sufficiently constrain N_{CO} and T_{CO} ; even for the targets with the highest quality Ly α and CO profiles (V4046 Sgr, DF Tau, and RECX 15), the uncertainty in the Ly α incident radiation causes a spread in the best-fitting N_{CO} and T_{CO} .

For fitting the (14 - 3) band we choose the regions between 1315 Å and 1322 Å carefully so as to avoid the 1316.55 and 1320.22 Å H₂ lines. We include the bandhead ($J'' < 5$), as it constrains N_{CO} for low T_{CO} . Also, the high J lines 1317 to 1320 Å must be included, as they constrain T_{CO} for a given value of N_{CO} . The range between 1316.5 and 1317.5 Å is generally excluded; interstellar H I absorption removes any knowledge of the Ly α flux which results in CO emission in this region. This corresponds to the $J'=6-9$, 9-12, and 12-15 transitions for the P, Q, and R branches, respectively. For the (14 - 4) band in V4046 Sgr, RECX 15, and DM Tau, we exclude flux between 1353.5 Å and 1354.5 for the same reason. Due to the computation time necessary to perform a grid search with our fluorescence model, we utilized a parameter optimization program starting from an initial coarse parameter grid.

Table 6 lists the CO fit parameters for all targets. Individual confidence intervals for each fit are smaller than the overall distribution of good fits. For this reason, the CO parameters we list are determined from the minimum and maximum parameter values, with the average of the two chosen as the mean value. Information for the (14 - 4) fits is also included for the three targets free of O I] emission (V4046 Sgr, RECX 15, and DM Tau). For targets with both (14 - 3) and (14 - 4) bands fits, the values of T_{CO} , N_{CO} , and Δv_{broad} agree within error bars. The required v_{CO} to fit each band differs by <10 km s⁻¹, which is below the accuracy of the COS wavelength solution. The lack of knowledge about the core of the Ly α profile is chiefly responsible for discrepancies between the model flux and that of the observed spectrum. Flux in the 1316 - 1317 Å region is typically under-predicted, in part due to the contaminating H₂ line, but also due to a more complicated absorbing outflow velocity distribution than we can reasonably model at present. Future work will include more complicated velocity profiles.

In DF Tau, RECX 15, and DM Tau, the high- J lines appear close to unresolved, which

help to strongly constrain Δv_{broad} . For DM Tau, HD 135344B, and DE Tau, the Ly α and/or CO S/N is low enough to make it difficult to constrain parameters as well as in the other targets. In RECX 11 the CO high- J lines are not resolved, although whether this is physical (due to velocity broadening of close-in CO) or merely due to the lower S/N in the band is not clear. As a result, Δv_{broad} has a wide range of values. DE Tau has a wide range of T_{CO} and N_{CO} , primarily due to its poorly-constrained Ly α . In GM Aur, HN Tau, DR Tau, DK Tau, and LkCa 15, the highly-absorbed Ly α and/or weak CO emission makes it impossible to constrain N_{CO} and T_{CO} better than several dex. Upper limits on N_{CO} cannot even be determined without better knowledge of their Ly α profiles. This will be possible in future work by partially reconstructing the Ly α profiles with measured H₂ fluxes. The models presented here assume isothermal CO. Relaxing this assumption would provide more flexibility in simultaneously fitting low- J and high- J emission lines. The COS observations of some TTS may have enough sensitivity to test disk models that predict a range of gas temperatures.

5. Discussion

5.1. Spatial Distribution of UV-CO

While the degeneracy of N_{CO} and the filling fraction prevents us from definitively determining the geometry of the UV-CO around the star, comparing our measured UV-CO properties with those of previously measured disk gas still reveals much about the spatial distribution of the gas. CO M-band fundamental emission ($v'=1, v''=0$) has been detected in several of these targets (Salyk et al. 2009; Najita et al. 2003; Bast et al. 2011; Pontoppidan et al. 2008). The gas in which the IR-CO emission originates is typically warm, thermally excited at temperatures as high as ~ 1500 K. Excitation diagrams of larger IR-CO samples can be adequately described by LTE models of a single-temperature slab of gas (Salyk et al. 2011b). Line-width analysis suggests that this emission is produced ≤ 1 AU from the central star. This information, combined with our measured UV-CO properties, presents two arguments that the UV-CO and IR-CO are spatially separate.

The UV-CO temperatures suggest separate populations, as our measured T_{CO} values are several times lower than typical IR-CO temperatures. The ro-vibrational levels of the CO ground electronic state are primarily populated based on the Einstein A coefficients of each v and J level, relative to the temperature-dependant collision rates and gas density. Models typically assume the IR-CO lies in a "skin" of the gas disk where densities are high enough for the rotational temperature to be equal to the kinetic temperature (Brittain et al. 2007). If this is also true for the UV-CO, then it must lie in a different region of the disk

where the gas kinetic temperature is lower. However, as temperature and density vary based on location in the disk it may be possible to find gas populations with thermal low- J levels, while the high- J levels are subthermal. In targets with well-defined high- J lines (V4046 Sgr, DF Tau, RECX 15, and DM Tau), these lines strongly influence the T_{CO} and N_{CO} fits. If densities are less than the critical densities for these high- J states (n_{crit} for $J=18$ at 400 K is $\sim 10^6$ cm $^{-3}$), then the derived temperature could be less than that determined from the thermalized lower states. In this case the CO density would have to be lower than the thermalized IR-CO, and therefore be in a different region of the disk.

Our measured Δv_{broad} values enable us to spatially constrain the UV-CO for the targets with well-defined high- J lines. Due to the blending of the UV-CO bands we cannot calculate the line width at zero intensity of individual J transitions, from which the inner CO radius is typically calculated (Brittain et al. 2009). However, assuming that the FWHM of the high- J lines is primarily due to orbital motion of the gas, we can calculate an average orbital radius R_{CO} :

$$R_{CO} = GM_* \left(\frac{2\sin(i)}{\Delta v_{broad}} \right)^2 \quad (2)$$

where G is the gravitational constant, M_* the stellar mass, and i the disk inclination, previously measured in the literature. For the values listed in Appendix A, we calculate R_{CO} for V4046 Sgr, DF Tau, RECX 15, and DM Tau: 2.76 AU, 2.16 AU, >3.04 AU, and >1.56 AU respectively. These radii are outside typical IR-CO radii ($R_{CO} < 1$ AU) calculated with Equation 2 and FWHM values from a survey of CTTSs from Salyk et al. (2011a). DF Tau is observed in their survey, with $R_{CO} = 0.15$ AU based on their CO FWHM (79 km s $^{-1}$), inclination (85°, Johns-Krull & Valenti (2001)), and M_* (0.27 M $_{\odot}$, Gullbring et al. (1998)). The Δv_{broad} values of RECX 15 and DM Tau are at the COS resolution, so we can only determine lower limits for R_{CO} in these disks. If thermal broadening, turbulent velocity, or disk winds contribute to the line widths, then R_{CO} could be even larger. Regardless, the temperature and Δv_{broad} disagreements lead us to the conclusion that the UV-CO and IR-CO are indeed separate populations.

In addition to the spatial differences between the two CO populations, there is no consistent agreement between the detection of IR-CO and UV-CO emission in our sample. DR Tau exhibits strong IR-CO emission (Salyk et al. 2008), but no UV-CO emission is observed. Conversely, DM Tau shows strong UV-CO emission, but no IR-CO emission. DF Tau, HD 135344 B, and DE Tau show both. Unfortunately, the stronger detection targets V4046 Sgr, RECX 15, and RECX 11 have yet to be observed in the M-band. Different morphology (UV-CO only, IR-CO only, both) is not a strong argument for separate gas populations. In

principle, different morphologies could occur in a single gas population if the Ly α pumping and IR excitation mechanisms are not physically related. Further study of a larger survey will be necessary to draw any solid conclusions on this matter, but the temperature and radius differences found in our analysis argue against a single gas population.

The UV-CO also appears to differ from H₂ previously studied in CTTSs. H₂ is best observed in the UV due to its lack of a permanent dipole moment. In the TW Hya disk, Herczeg et al. (2004) detected H₂ emission, measuring $T_{H_2} = 2500^{+700}_{-500}$ K and $\log N_{H_2} = 18.5^{+1.2}_{-0.8}$ cm⁻². In addition, France et al. (2011, in press) observed H₂ in absorption in the spectrum of AA Tau, measuring $T_{H_2} = 2500^{+800}_{-700}$ K and $\log N_{H_2} = 17.9^{+0.6}_{-0.3}$ cm⁻². As with the IR-CO, the differences in temperature suggest that the two populations reside in different regions of the disk. Similar to the conclusion of France et al. (2011, in press) about the AA Tau CO and H₂, a subthermal CO population could explain the disparity. If the CO and H₂ are indeed cospatial, then the similarity of N_{CO} and N_{H_2} derived from UV spectra would seem to violate the CO/H₂ ratio of 10⁻⁴ (Lacy et al. 1994). It is possible that during the process of star and planet formation, the CO is enhanced and/or H₂ removed from the gas phase in a way that increases the CO/H₂ ratio. As suggested by France et al. (2011a), collisions of large grains or planetesimals may re-release gas-phase CO into the disk. In addition, H₂ could be selectively dissociated by radiation in the inner disk. However, in order to truly compare both molecules the UV H₂ and CO emission must be studied with a self-consistent approach.

Comparing N_{CO} and T_{CO} with other disk properties published in the literature does not yield any significant correlations. For example, an increase in rotational temperature vs. accretion rate might be expected; higher accretion rates create greater FUV continuum and Ly α emission, which in turn heats the gas to higher temperatures via photo-electric heating. The FUV photons could also populate higher J states in a way not possible by collisional processes alone. However, as seen in Figure 6, there appears to be no correlation between temperature and accretion rate. Over more than two orders of magnitude of accretion rate, targets with strong CO and Ly α emission well constrain the temperature between 200 K and 500 K with no apparent trend. Only in DM Tau and DE Tau is the temperature less constrained, but this is primarily due to their highly absorbed Ly α profiles.

5.2. CO Band Flux and Disk Evolution

The total CO emission from our sample can be evaluated by integrating the (14 - 3) band flux. Referring to Figures 1 and 2, the (14 - 3) bands are summed between the dashed grey lines at 1315.2 Å and 1320 Å. The continuum level $C_{(14-3)}$ is determined from the surrounding

line-free spectrum, and denoted by the horizontal solid black/blue/red line. The blue and red lines identify the low and high J state emission regions, which are separately integrated. The dashed orange lines above and below mark the $1-\sigma$ error of the continuum level. The integrated flux values are listed in Table 7. For DR Tau and LkCa 15 we list zero flux, which actually corresponds to an overall negative flux. Shown in Figure 7 is ratio of Ly α blue to red flux vs. the ratio of CO blue (low- J) to red (high- J) flux. This demonstrates our initial correlation between the relative strength of the Ly α photo-exciting radiation and the resulting CO emission.

To test correlations with disk evolutionary properties we use the total CO ($14 - v''$) luminosity, calculated from the ($14 - 3$) band fluxes:

$$L(CO)_{14-3} = \frac{F(14-3)}{B(14-3)} 4\pi d^2 \quad (3)$$

where $F(14 - 3)$ is the integrated band flux, $B(14 - 3)$ is the branching ratio listed in Table 3, and d is the distance to the target. In Figure 8, $L(CO)_{14-3}$ shows a strong correlation with mass accretion rate. This is not surprising, as the amount of photo-exciting Ly α is expected to increase with accretion. Combined with the weak opacity of the CO ($14 - 0$) band, we are simply seeing the re-emission of most of the Ly α photons capable of being absorbed.

As a result, \dot{M}_{acc} cannot be used as a proper evolutionary diagnostic, despite the fact that accretion does decrease as a CTTS ages (Hartmann et al. 1998; Fang et al. 2009). We attempt to correct for this by dividing the CO luminosity by \dot{M}_{acc} to compare with age. Figure 9 shows this relation, with a trend towards more CO emission in more evolved systems. One interpretation is that as inner disks clear, the Ly α can more readily illuminate the CO. V4046 Sgr, HD135344B, DM Tau, and GM Aur are all confirmed transitional objects with cleared dust holes in the center of their disks (Jensen & Mathieu 1997; Rodriguez et al. 2010; Brown et al. 2009; Calvet et al. 2005), which would allow the Ly α photons to reach more CO. CO-IR emission from HD 135344B and GM Aur support this argument, as the CO-IR appears inside the inner dust radii (Pontoppidan et al. 2008). However, no UV-CO is detected in the transitional disk LkCa 15. A second interpretation would be the increase of gas-phase CO through grain and planetesimal collisions (France et al. 2011a). CO produced in the dust disk could be transported to the inner disk surface as the system evolves.

5.3. C IV Photoexcitation

Finally, we measure flux in the 1712 - 1720 Å region, assuming that emission in this region is due to the (14 - 12) band. Of particular note is how large these fluxes are compared to what would be expected from simply applying branching ratios to the (14 - 3) fluxes. Figure 10 shows this graphically; for one of the best-fit (14 - 3) models of RECX 11, the corresponding (14 - 12) emission is plotted over its observed 1712 Å emission. This trend is seen significantly in all targets and also shown in the bottom of Figure 10 with the dashed line representing the expected (14 - 12) integrated flux for a given (14 - 3) integrated flux. If this emission is due to CO, some process other than Ly α photo-excitation must be responsible. It is likely the result of C IV photo-excitation, which is capable of exciting the $v'=0$ level of CO (France et al. 2011a). The resulting (0 - 3) transition creates a band of emission at 1712 - 1720 Å. Cursory analysis shows that there is sufficient flux in the observed C IV lines to produce these levels of emission.

6. Conclusions

In this paper, we have characterized UV photo-excited CO emission in a sample of T Tauri star spectra from the COS instrument on *HST*, reaching the following conclusions:

1. Electronic transitions of CO are observed in approximately half of the T Tauri stars surveyed, primarily due to photo-excitation by Ly α emission coincident with the $A^1\Pi - X^1\Sigma^+$ (14 - 0) absorption band. The resulting transitions from the $v' = 14$ state appear as emission bands between 1280 Å and 1712 Å. The (14 - 3) and (14 - 4) bands are most conducive to spectral modeling.
2. The observed Ly α emission is not sufficient to reproduce the observed CO band emission; thus, we create a model to simulate the local Ly α radiation field incident on the CO molecules. With two Gaussian emission components created at the star, followed by outflow and ISM absorbers, we are able to recreate the observed Ly α profiles and CO band emission.
3. For seven of the twelve targets, we measure CO temperatures between \sim 100 and 1500 K, and CO column densities of $\sim 10^{18} - 10^{19}$ cm $^{-2}$. Both strong Ly α and CO emission are required in order to constrain N_{CO} to within 0.5 dex; for these 5 targets temperatures are below 600 K. The high- J lines in several targets are narrow enough to constrain the CO to outside 1 AU.
4. There are no strong correlations between our calculated N_{CO} and T_{CO} values and

other previously determined parameters. Our measured CO temperatures are lower than previously measured IR-CO and UV-H₂ in CTTSs. The temperature differences, along with our calculated emission line widths, imply that the UV-CO and IR-CO are different populations.

5. If the UV-CO and UV-H₂ are co-spatial, our N_{CO} values are 3-4 orders of magnitude greater than those expected from previously measured H₂ abundances and dense cloud CO/H₂ ratios. Further analysis of the H₂ lines, as well as more detailed geometric modeling, is essential to understand the relationship between the UV-CO, IR-CO and H₂.

6. CO emission increases with \dot{M}_{acc} because of the corresponding increase in the Ly α radiation field. Upon correcting for \dot{M}_{acc} , there is still a correlation between CO emission and age. We conclude that as CTTSs evolve and disk dust and H I clears, the Ly α photons are able to reach more CO.

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A. Targets

A.1. η Chamaeleontis Targets

The η Chamaeleontis cluster is located at a distance of 97 pc. Age estimates of its constituent stars range from 4 - 10 Myr (Lawson et al. 2001). It is unusual compared to other star forming regions; Sicilia-Aguilar et al. (2009) measured a disk fraction of $\sim 50\%$, much higher than expected for an older region. There is also a higher percentage of disk systems in transition objects than found in other clusters (Megeath et al. 2005).

A.1.1. RECX 11

RECX 11 is a weakly accreting CTTS. Its low accretion rate ($\dot{M} = 4 \times 10^{-11} M_{\odot} yr^{-1}$) would imply that photoevaporation should have cleared the inner disk, but this is not seen (Gautier et al. 2008). UV-H₂ emission implies the presence of gas (France et al. 2011b). Ingleby et al. (2011) also find low accretion values ($\dot{M} = 3 \times 10^{-10} M_{\odot} yr^{-1}$) using the same data set presented here and argue that negligible planet formation has occurred during the

disk evolution. From H₂ line profiles, they measure rotational velocities fast enough to place the molecular gas as close as 3 R_{*}.

A.1.2. RECX 15

RECX 15 is an active accretor; the H α EW is 90.0 Å, and its profile is indicative of a blue-shifted wind component (Lawson et al. 2004). The SED implies a continuous disk from close to the stellar surface out to tens of AU or more, where the disk flares strongly (Gautier et al. 2008; Woitke et al. 2011). However, the H₂ (1 - 0) S(1) emission line width measured by Ramsay Howat & Greaves (2007) places the UV-H₂ a few AU from the star. RECX 15 has a stellar mass of 0.2 M_⊙ and inclination angle of 40° (Woitke et al. 2011).

A.2. Taurus-Auriga Targets

The Taurus-Auriga complex is one of the nearest star forming regions at a distance of ∼140 pc. Its paucity of O/B stars, combined with a low interstellar radiation field, enable the persistence of many disks (Luhman et al. 2010). Disk fractions in the Taurus region have been measured from ∼ 45% for low-mass objects (M=0.01-0.3 M_⊙) to ∼75% for solar-mass stars. Some of these objects have been studied in exceptional detail, while others have only been measured in large surveys.

A.2.1. DM Tau

DM Tau is considered to be a transitional disk due to its flux deficit shortward of 8 μm (Calvet et al. 2005). Resolved images and spectra of dust continuum emission show an inner hole of 19 AU with little dust inside this region (0.0007 lunar masses), although with a detectable accretion rate ($\sim 1.3 \times 10^{-8} M_{\odot} \text{yr}^{-1}$) there must be gas present (Calvet et al. 2005; Andrews et al. 2011). The disk exhibits a rich chemistry; organic molecules CO, CN, CS, HCN, HNC, HCO⁺, C₂H, and N₂H⁺ have been measured in emission (Dutrey et al. 1997). The abundances of these molecules show that stellar UV photons have strong influence over their creation and destruction. DM Tau has a stellar mass of 0.43 M_⊙ and inclination angle of 35° (Hartmann et al. 1998; Andrews et al. 2011).

A.2.2. *HN Tau*

HN Tau is a binary system (0.7 and $0.4 M_{\odot}$) with strong jets that expand with distance from the primary star (Hartigan et al. 2004). In this work we study only the spectrum of HN Tau A, due to its separation from HN Tau B ($3.1''$). M-band observations from NIRSPEC on Keck reveal weak CO emission, but strong absorption (J.M. Brown, private communication). While its exact inclination is not known, France et al. (2011a) measure CO electronic transitions detected in absorption, which could be indicative of disk gas viewed at high inclination.

A.2.3. *DR Tau*

DR Tau is a K4 star with an accretion rate that appears to vary by over two orders of magnitude (Gullbring et al. 2000; Johns-Krull & Gafford 2002). DR Tau shows strong molecular emission at IR wavelengths. Rotational transition emission lines of H_2O are seen, in high J_s (20-50), as well as CO and OH at temperatures around 1000 K (Salyk et al. 2008). By modeling the gas as an isothermal LTE ring, the emission radii are 1.0, 0.8, and 0.7 AU for H_2O , CO, and OH, respectively. Time-lapsed CO measurements show variability between overtone emission and absorption (Lorenzetti et al. 2009).

A.2.4. *DK Tau*

DK Tau is a binary system consisting of a K9 and M0 with ~ 350 AU separation. From mm imaging both appear to have individual circumstellar disks (Jensen & Akeson 2003), however our observation was targeted on DK Tau A. Based on the UV Unlike the majority of binary systems in which both stars contain disks, the A and B disks may be misaligned (Kraus & Hillenbrand 2009). Some molecular emission has been detected in the DK Tau disks, particularly sub-mm ^{13}CO emission (Greaves 2005).

A.2.5. *DF Tau*

A binary pair (M1 + M3.5) with 13 AU separation, DF Tau may be as young as 0.1 Myr (Ghez et al. 1997). Our spectrum samples the primary, as the secondary is very dim at UV wavelengths. Strong accretion has been measured from a UV continuum (Herczeg et al. 2006; France et al. 2011b), as have UV H_2 emission lines. H_2 appears both in Ly α /C IV

pumped emission, and in absorption in an outflow (Yang et al. 2011). DF Tau displays a rich spectrum of IR-CO emission lines; the CO (1 - 0) R(3) shows a FWHM of $\sim 74 \text{ km s}^{-1}$, implying a close orbit to the star (Najita et al. 2003). Greaves (2005) also detect sub-mm emission from ^{13}CO . DF Tau has a mass of $0.27 M_{\odot}$ and inclination of 85° (Gullbring et al. 1998; Johns-Krull & Valenti 2001)

A.2.6. LkCa15

LkCa15 is a pre-transitional CTTS with a ringed gap in its disk, likely caused by a (proto)planet (Kraus & Ireland 2011). From mid-IR imaging the inner rim of the optically thick dust disk is calculated to be 46 AU, with some remnant dust inside (Espaillat et al. 2007; Piétu et al. 2006). Najita et al. (2003) measured fundamental CO emission as close as 0.1 AU.

A.2.7. GM Aur

GM Aur is a massive transitional disk; the *Spitzer* Infrared Spectrograph (IRS) SED is best explained by 0.02 lunar masses of dust within a 5 AU inner disk, and an outer disk with a 24 AU inner radius (Calvet et al. 2005). High resolution sub-mm observations support this picture, finding a 28 AU hole in the optically thick dust disk (Andrews et al. 2011). Salyk et al. (2007) measured near-IR CO emission within $0.2_{-0.2}^{+0.4}$ AU, at high rotational temperatures, coincident with the measured dust inner rim. Ne II was measured in the surrounding disk at the systemic velocity of 14.8 km s^{-1} (Najita et al. 2009). Along with a profile broadened by rotational velocity, this would suggest the gas lies in the disk.

A.3. Isolated Targets

Two of the targets do not reside in star-forming regions. V4046 Sagittarius is a ~ 9 Myr old CTTS, a closely orbiting ($\sim 9 R_{\odot}$ separation) binary with combined mass of $1.78 M_{\odot}$ (Stempels & Gahm 2004) at a distance of 83 pc (Rodriguez et al. 2010). At an inclination of 36° , the disk appears to be coplanar with the central binary stars (Monin et al. 2007). Kastner et al. (2008) measured molecular mm emission lines from CO, HCN, CN, and HCO^{+} ; the double peaked CO profiles confirm outer disk Keplerian rotation. Rodriguez et al. (2010) imaged the system in $^{12}\text{CO}(2-1)$ and $^{13}\text{CO}(2-1)$ millimeter emission, measuring the molecular disk extent to ~ 370 AU and mass of ~ 110 Earth masses. HD 135344B is an F star in a

visual binary, and has a disk with a gap from .45 to 45 AU (Brown et al. 2007). It is more massive than the other stars in our sample. CO has been detected in emission inside the hole out to \sim 15 AU (Brown et al. 2009). The distribution of gas column densities agree with models of a Jovian planet orbiting between 10 and 20 AU (Pontoppidan et al. 2008).

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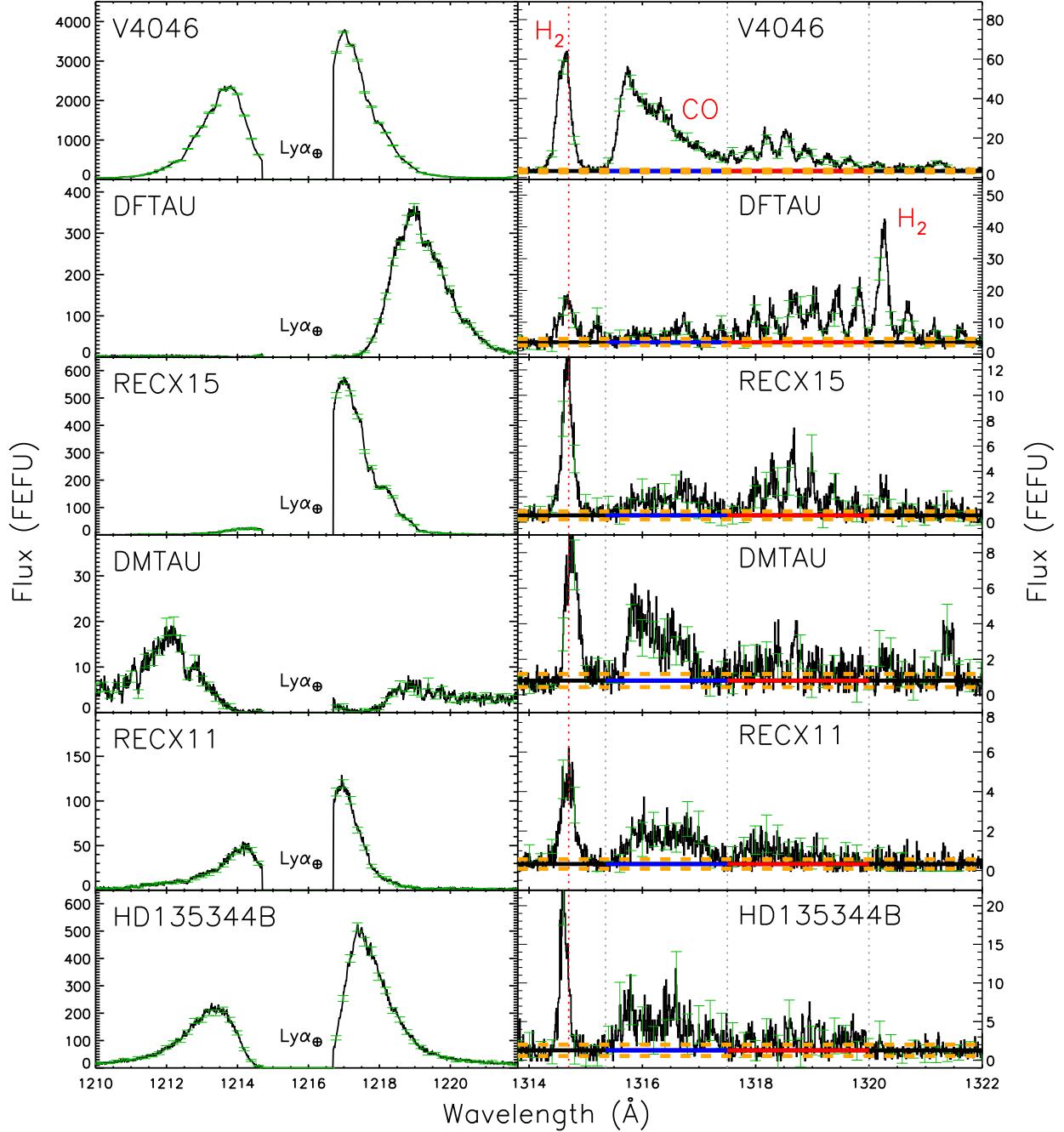


Fig. 1.— Extinction-corrected Ly α profiles (left column) with geocoronal Ly α removed, and CO (14 - 3) emission bands (right) for each of the sample targets. In the CO bands, an H_2 line is identified with a vertical dotted red line, and the dotted grey lines bound the high and low- J sides of the (14 - 3) band. The horizontal black/blue/red line identifies the local continuum, with the dashed orange lines at the 1σ error in the continuum average. Representative error bars are shown in green.

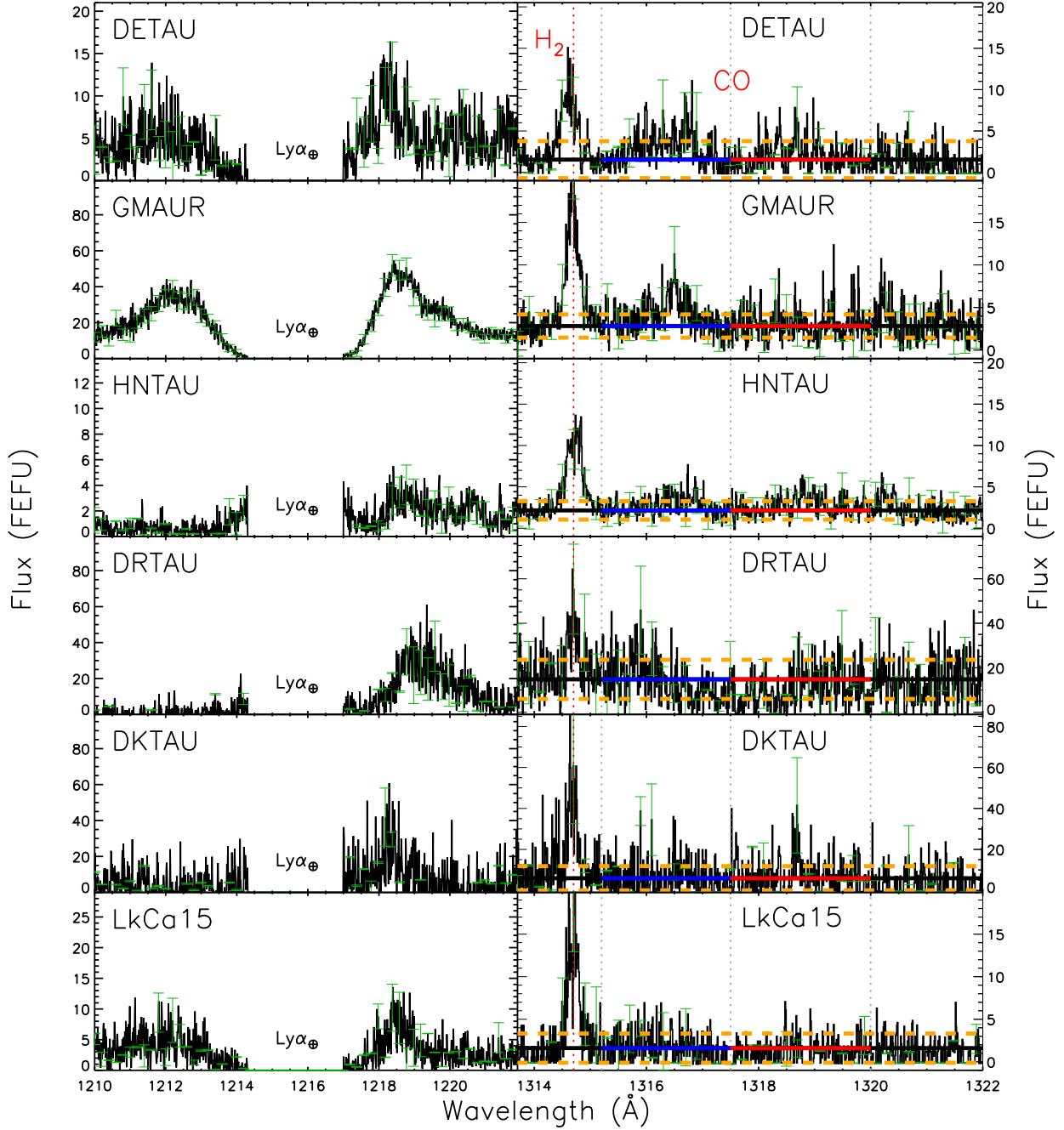


Fig. 2.— Extinction-corrected Ly α profiles (left column) with geocoronal Ly α removed, and CO (14 - 3) emission bands (right) for each of the sample targets. In CO bands, an H₂ line is identified with a vertical dotted red line, and the dotted grey lines bound the high and low- J sides of the (14 - 3) band. The horizontal black/blue/red line identifies the local continuum, with the dashed orange lines at the 1 σ error in the continuum average. Representative error bars are shown in green. DR Tau and LkCa 15 are non-detections.

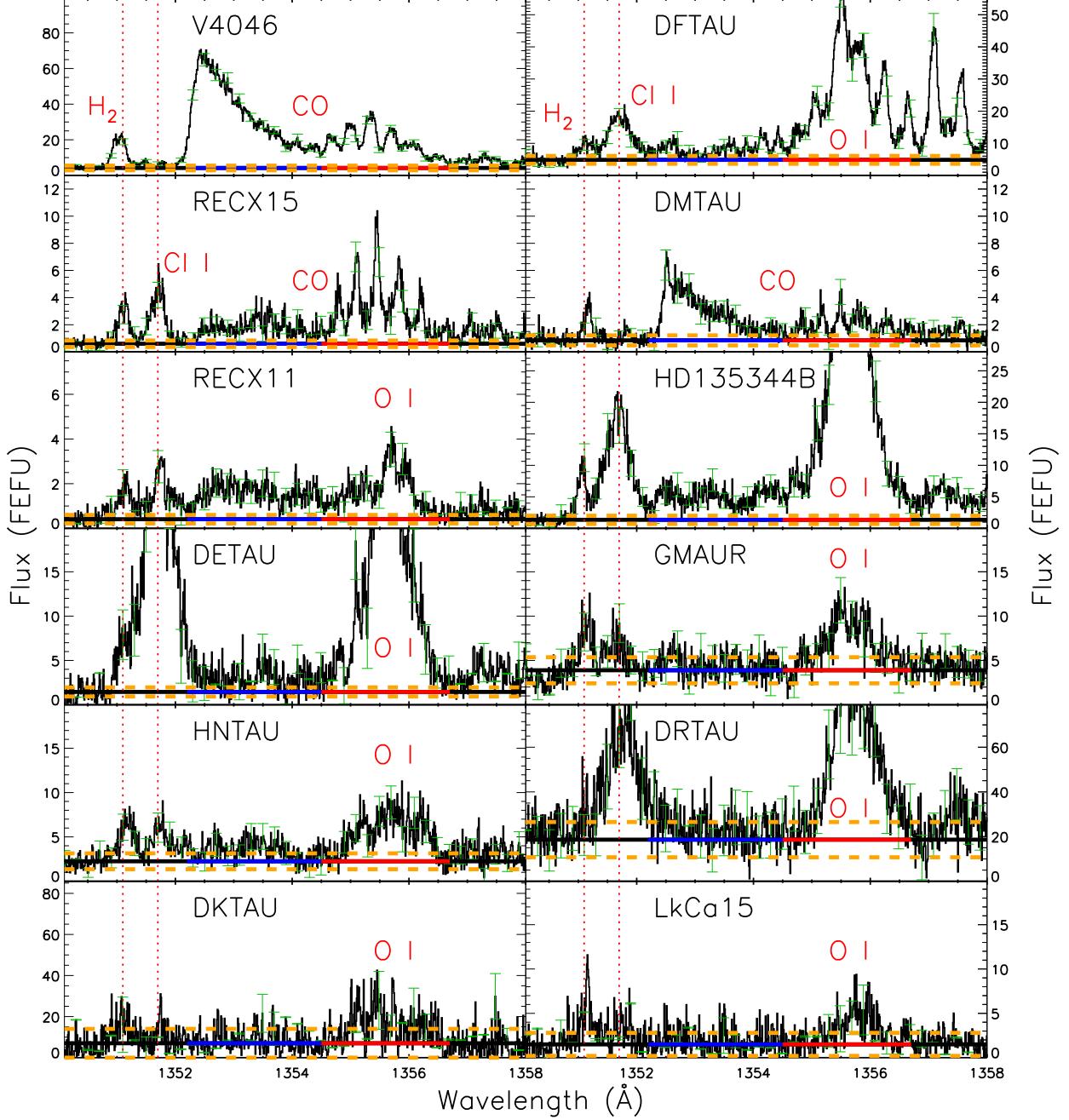


Fig. 3.— (14 - 4) transitions for targets as identified. The horizontal lines show continuum levels, with dashed orange lines indicating the upper and lower limits for summation. The solid blue and red lines mark the wavelength range for the high and low- J regions as described in the text; vertical gray lines bound these regions. Vertical dashed red lines mark contaminating lines. Representative error bars are shown in green. H_2 and $Cl\ I$ emission lines are identified with red-dotted lines, and $O\ I$ emission features are identified in the spectra in which they appear.

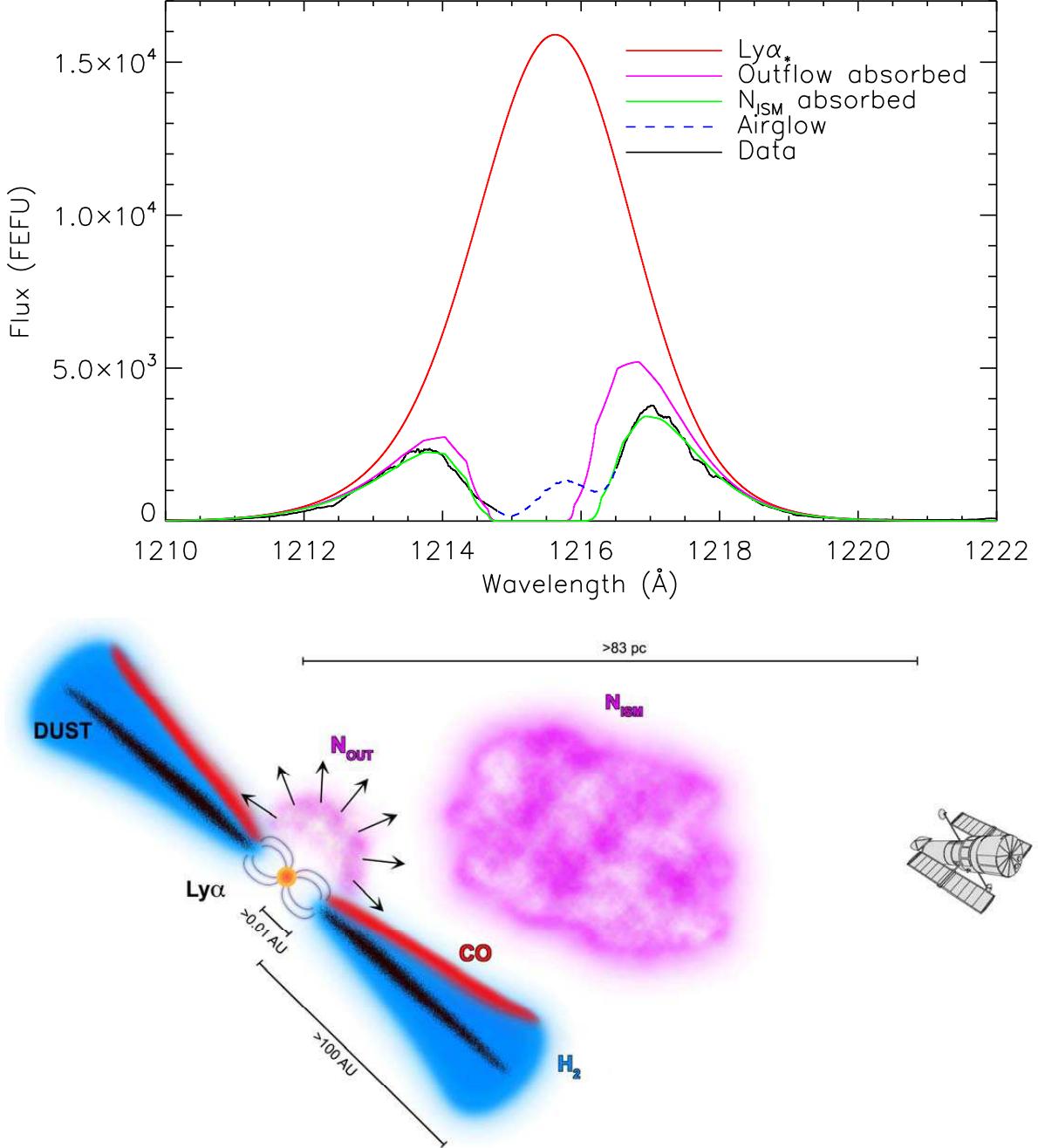


Fig. 4.— V4046 Sgr Ly α data with overlaid model profiles (top), and drawing of disk that produces the flux we see (bottom). Ly α emission (red profile) is created near the star, then absorbed by an outflow (magenta profile), and additionally absorbed by interstellar H I (green profile) to produce the observed data. The CO in the disk sees the outflow-absorbed emission. Disk height is not to scale relative to extent.

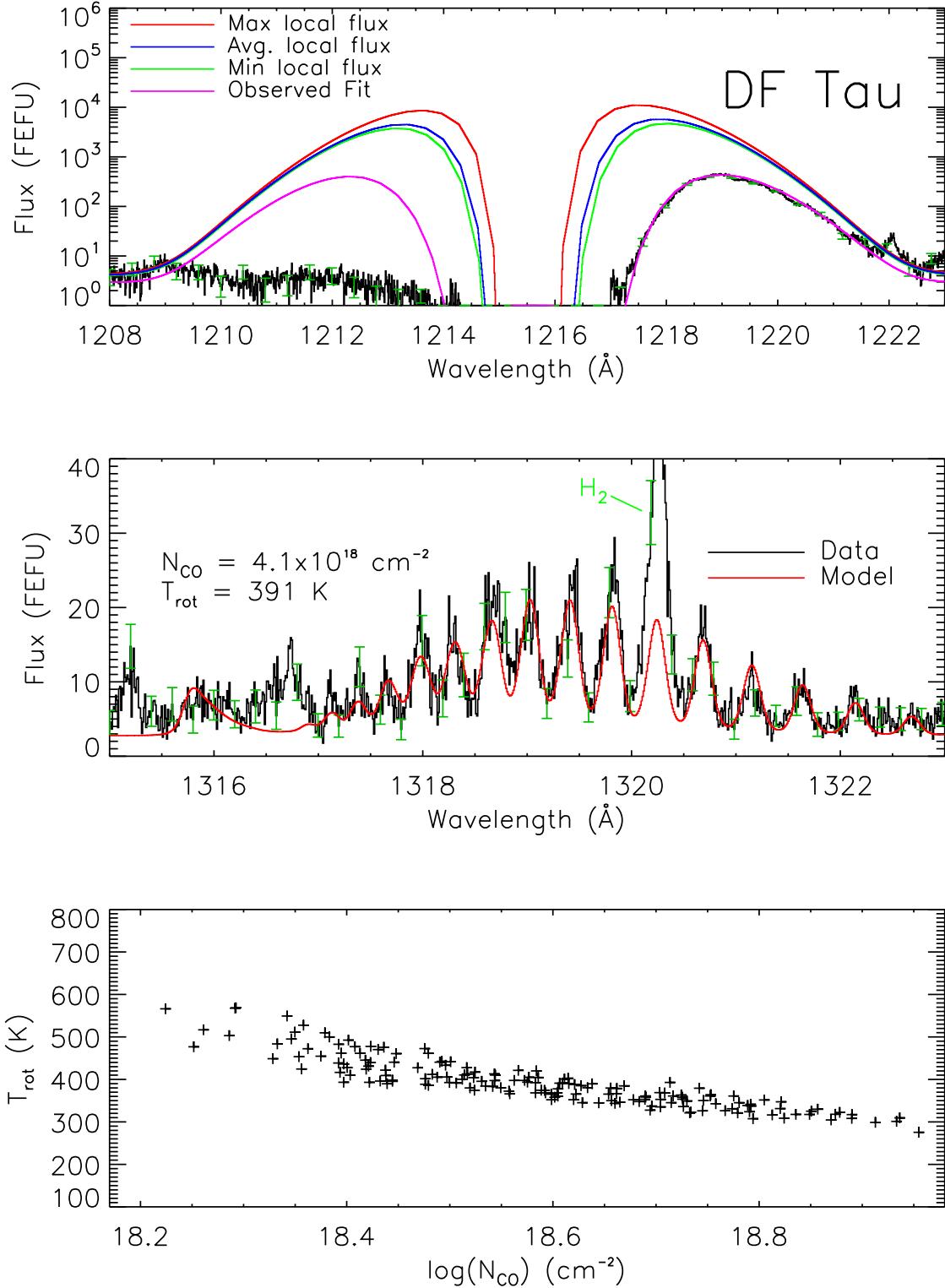


Fig. 5.— DF Tau model results. Top: Ly α data overplotted with minimum, average, and maximum model Ly α profiles. Middle: CO (14 - 3) band with model from average T_{CO} and N_{CO} . Bottom: T_{CO} vs. N_{CO} for best fit models. The red square shows the values used in the middle plot.

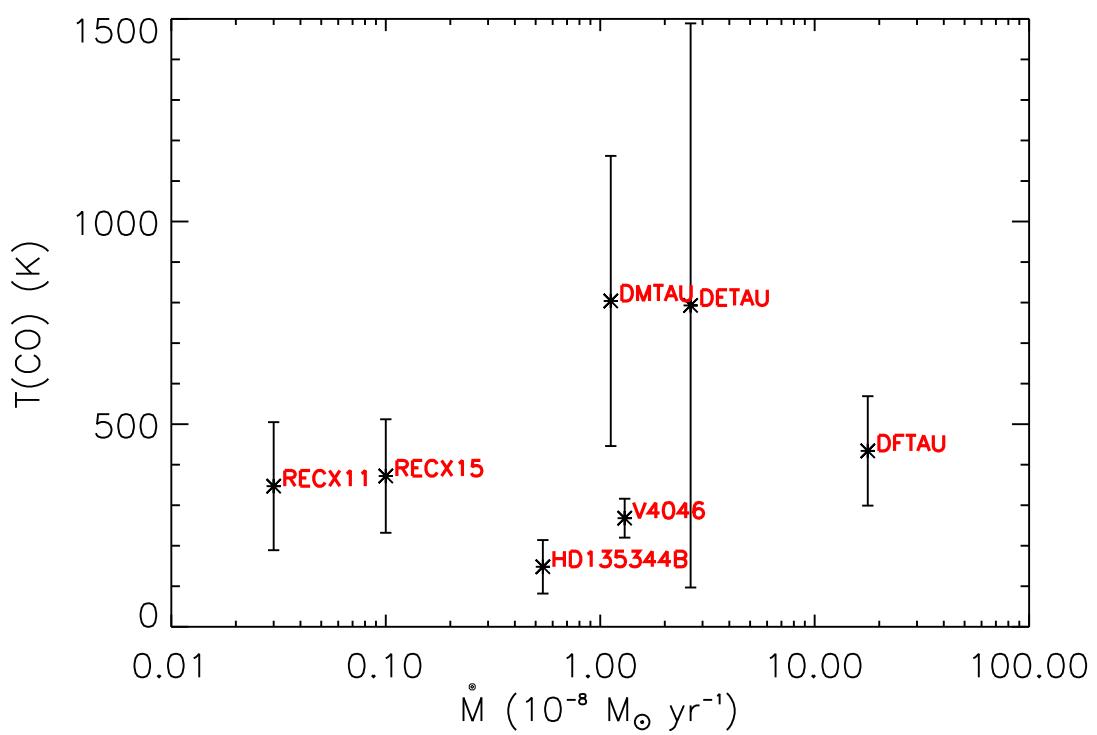


Fig. 6.— Model-derived T_{CO} vs. literature accretion rates.

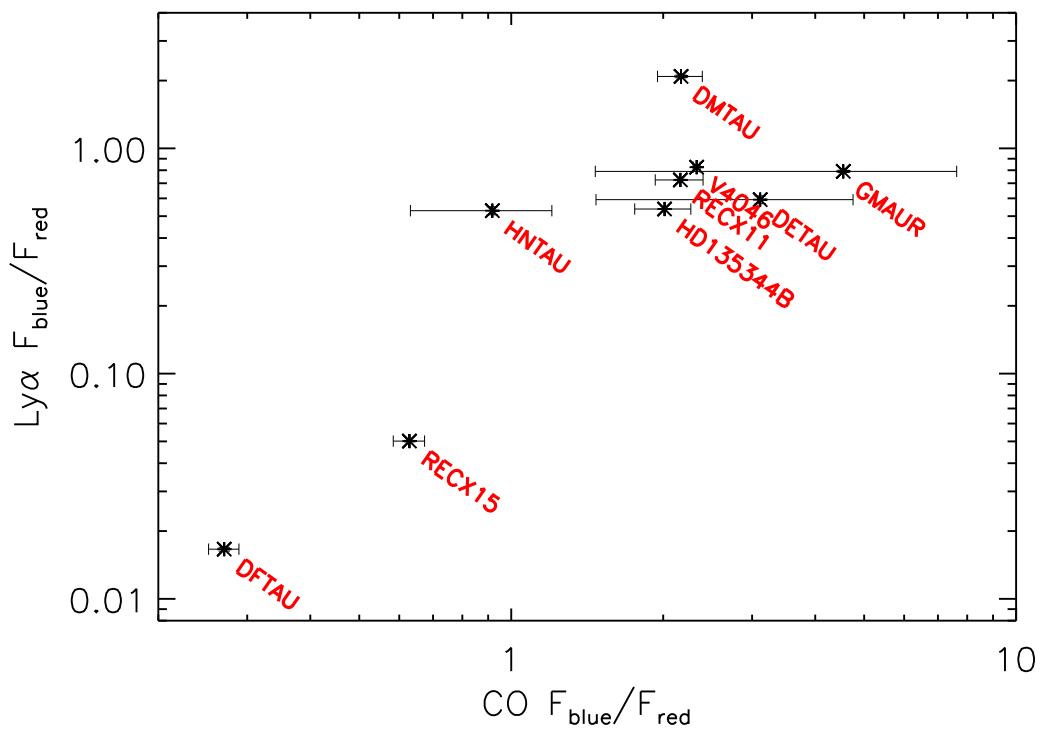


Fig. 7.— Ratios of blue to red flux for Ly α and CO. The blue and red Ly α fluxes were summed from 1210 - 1214.7 Å and 1216.7 - 1222 Å, respectively.

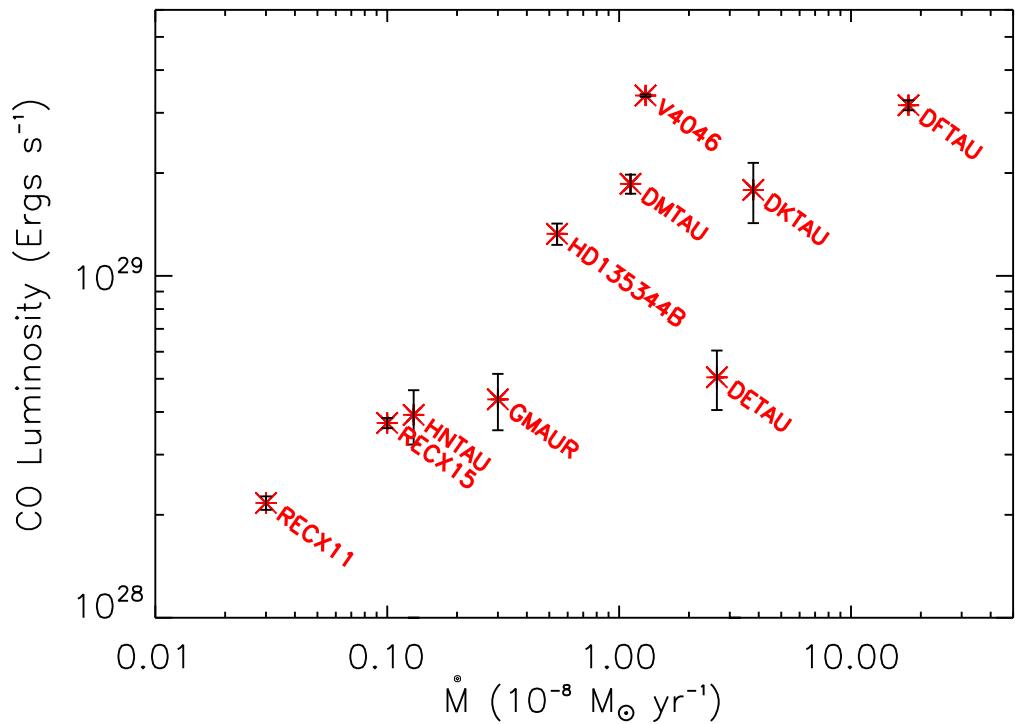


Fig. 8.— CO (14 - v'') band luminosity vs. literature accretion rates for positive detection targets. The CO luminosities were calculated by applying branching ratios to the (14 - 3) luminosities.

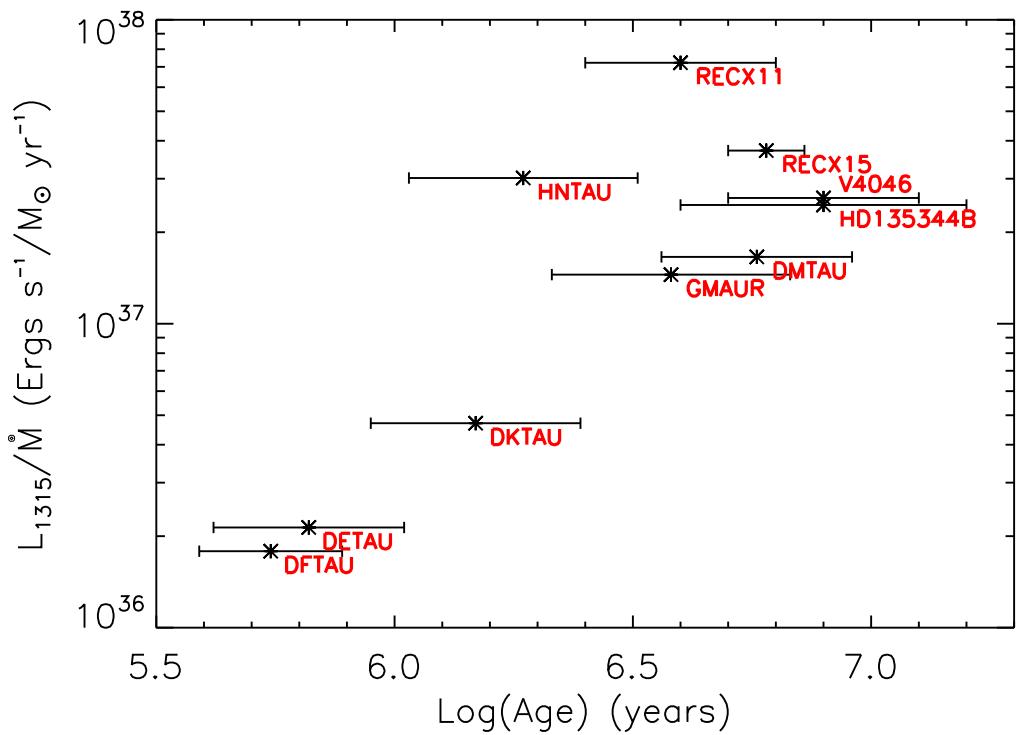


Fig. 9.— CO (14 - ν'') band luminosities, divided by literature accretion rates, demonstrate an increasing correlation with age.

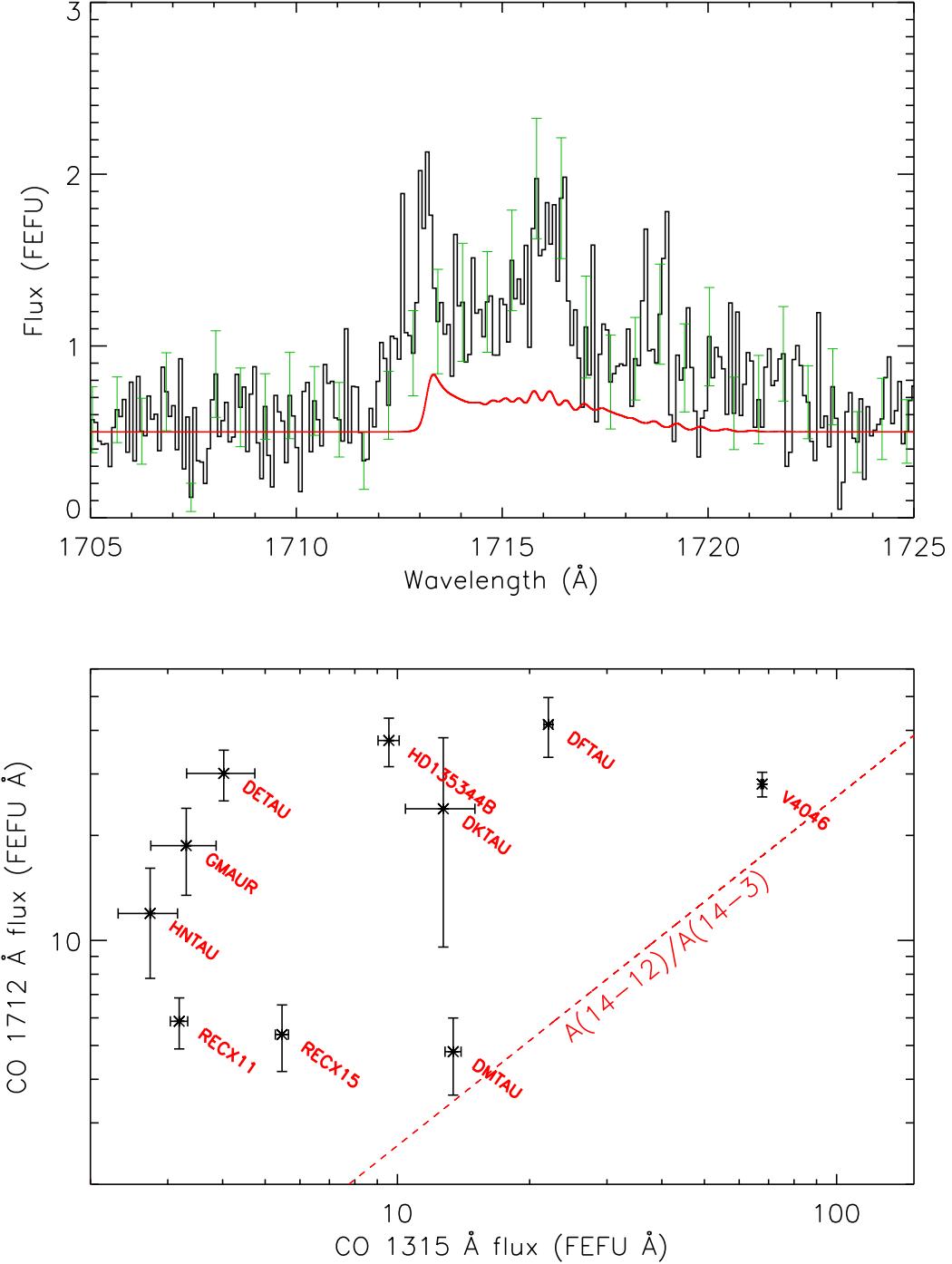


Fig. 10.— Top: RECX 11 (14 - 12) model emission (red) over the observed emission. The data have been smoothed to one spectral resolution element for display purposes. Summed fluxes for the 1712 Å emission vs. the 1316 Å band, for each target. The dashed red line shows the expected (14 - 12) flux expected from branching ratios for a given (14 - 3) flux.

Table 1: *HST-COS* observing log

Object Object	R.A. (J2000)	Dec. (J2000)	Date	T_{exp} (s) (G130M,G160M)	Program
RECX 11	08 ^h 47 ^m 01.28 ^s	-78° 59' 34.1"	2009 December 12	3645, 4514	11616
DF Tau	04 ^h 27 ^m 02.81 ^s	+25° 42' 22.1"	2010 January 11	4828, 5549	11533
DK Tau	04 ^h 30 ^m 44.25 ^s	+26° 01' 24.5"	2010 February 4	1585, 1971	11616
LkCa15	04 ^h 39 ^m 17.79 ^s	+22° 21' 03.1"	2010 February 5	1712, 1725	11616
RECX 15	08 ^h 43 ^m 18.43 ^s	-79° 05' 17.7"	2010 February 5	3891, 4502	11616
DR Tau	04 ^h 47 ^m 06.22 ^s	+16° 58' 42.6"	2010 February 10	1704, 1746	11616
HN Tau	04 ^h 33 ^m 39.37 ^s	+17° 51' 52.1"	2010 February 10	5725, 4529	11616
HD135344B	15 ^h 15 ^m 48.42 ^s	-37° 09' 16.3"	2010 March 14	2666, 4836	11828
V4046 Sgr	18 ^h 14 ^m 10.49 ^s	-32° 47' 34.2"	2010 April 27	4504, 5581	11533
GM Aur	04 ^h 55 ^m 10.98 ^s	+30° 21' 59.1"	2010 August 19	2128, 1861	11616
DE Tau	04 ^h 21 ^m 55.69 ^s	+27° 55' 06.1"	2010 August 20	2068, 1851	11616
DM Tau	04 ^h 33 ^m 48.74 ^s	+18 ° 10' 09.7"	2010 August 22	3459, 3770	11616

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Table 2: Target Parameters

Object	Sp. Type	Group	d (pc)	A _V	\dot{M} ($10^{-8} M_{\odot}$ yr $^{-1}$)	$\log_{10}(\text{Age})$ (yrs)	References
V4046 Sgr	K5, K7	Isolated	83	0	1.3	7.08 ± 0.20	6,8,10
DF Tau	M1, M3.5	Taurus-Aur	140	0.5	17.69	5.74 ± 0.15	1, 11
RECX 15	M3.4	η Cha	97	0	0.1	6.78 ± 0.08	5,12
DM Tau	M1	Taurus-Aur	140	0.5	1.12	6.76 ± 0.20	3,7,11
RECX 11	K6.5	η Cha	97	0	0.03	6.60 ± 0.20	5,13
HD135344B	F4	Isolated	140	0.3	0.54	6.90 ± 0.30	4,14
DE Tau	M0	Taurus-Aur	140	0.62	2.64	5.82 ± 0.20	1,11
GM Aur	K7	Taurus-Aur	140	0.96	0.3	6.58 ± 0.25	1,11
HN Tau	K5, M4	Taurus-Aur	140	0.65	0.13	6.27 ± 0.24	1,11
DR Tau	K4	Taurus-Aur	140	1.2	3.16	5.29 ± 0.34	2,9,11
DK Tau	K9, M0	Taurus-Aur	140	1.42	3.79	6.17 ± 0.22	1,11
LkCa15	K5	Taurus-Aur	140	.62	0.13	6.35 ± 0.26	3,7,11

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(1) Gullbring et al. (1998); (2) Gullbring et al. (2000); (3) Kenyon & Hartmann (1995); (4) Garcia Lopez et al. (2006); (5) Lawson et al. (2004); (6) Stempels & Gahm (2004); (7) Hartmann et al. (1998); (8) France et al. (2011b); (9) White & Ghez (2001); (10) Quast et al. (2000); (11) Kraus & Hillenbrand (2009); (12) Ramsay Howat & Greaves (2007); (13) Lawson et al. (2001); (14) van Boekel et al. (2005)

Table 3: CO Fourth Positive bands identified in sample CTTSs.

Band ID	λ_{obs} (Å) ^a	Pumping Source	λ_{pump} (Å)	$A(v' - v'')$ (s^{-1}) ^b	$B(v' - v'')$ ^c
(14 - 2)	1280.5	H I Ly α	1214 - 1222	4.0×10^8	0.04
(14 - 3)	1315.7	H I Ly α	1214 - 1222	1.5×10^9	0.17
(14 - 4)	1352.4	H I Ly α	1214 - 1222	2.7×10^9	0.30
(14 - 5)	1390.7	H I Ly α	1214 - 1222	1.6×10^9	0.17
(14 - 7)	1472.6	H I Ly α	1214 - 1222	9.5×10^8	0.10
(14 - 8)	1516.3	H I Ly α	1214 - 1222	5.1×10^8	0.06
(14 - 10)	1610.1	H I Ly α	1214 - 1222	5.6×10^8	0.06
(14 - 12)	1713.2	H I Ly α	1214 - 1222	3.9×10^8	0.04
(0 - 1)	1597.3	C IV	1544 - 1551	4.3×10^9	0.33
(0 - 2)	1653.2	C IV	1544 - 1551	3.6×10^9	0.27
(0 - 3)	1712.4	C IV	1544 - 1551	1.8×10^9	0.14

^aCO wavelengths are taken from Kurucz (1993).

^bSummed A values for each ($v' - v''$) band.

^cBranching ratios for each vibrational band, relative to all ($v' - v''$) bands.

Table 4: Best-fit Ly α parameters

Object	I ₁ (0) (FEFU)	σ_1 (km s ⁻¹)	v ₂ (km s ⁻¹)	I ₂ (0) (FEFU)	σ_2 (km s ⁻¹)	C _{Lyα} (FEFU)
V4046 Sgr	10122 ± 3531	263 ± 40	-72 ± 35	5232 ± 2840	368 ± 31	7.0 ± 3.0
DF Tau	52084 ± 943	374 ± 0	-	-	-	4.5
RECX 15	16 ⁺²⁹ ₋₁₆	265 ⁺²⁹² ₋₂₆₅	208 ± 28	906 ± 146	224 ± 10	0.8 ± 0.2
DM Tau	18 ⁺³⁴ ₋₁₈	339 ⁺³⁶⁵ ₋₃₃₉	-133 ± 20	270 ± 64	437 ± 14	4.0
RECX 11	1376 ± 1	197 ± 0	-	-	-	2.0
HD 135344B	2487 ± 1109	346 ± 52	118 ± 151	2293 ± 2168	158 ± 95	0.8 ± 0.2
DE Tau	4835 ± 4590	249 ± 59	-13 ± 78	1693 ⁺¹⁹¹⁴ ₋₁₆₉₃	225 ± 152	1.1 ± 0.5

Columns: (2) Peak intensity of rest gaussian emission component; (3) Width parameter for rest gaussian emission component; (4) Peak intensity of secondary gaussian emission component; (5) Width parameter for secondary gaussian emission component; (6) Local continuum for Ly α

Table 5: Best-fit Ly α parameters (continued)

Object	v_{out} (km s $^{-1}$)	N_{out} (10 $^{19}cm^{-2}$)	N_{ISM} (10 $^{19}cm^{-2}$)	χ^2_ν
V4046 Sgr	-96 ± 15	3.0 ± 0.8	1.1 ± 0.4	104 - 136
DF Tau	-36 ± 7	14.0 ± 1.5	49.0 ± 1.4	< 1.08
RECX 15	-85 ± 36	1.1 ± 1.0	0.9 ± 0.2	2.2 - 20
DM Tau	119 ± 178	1.4 ± 1.2	41.8 ± 2.1	< 1.16
RECX 11	-99 ± 4	1.9 ± 0.1	3.3 ± 0.1	< 13
HD 135344B	-99 ± 26	5.6 ± 2.7	7.1 ± 1.6	0.76 - 7.1
DE Tau	-103 ± 56	11.4 ± 7.7	34.8 ± 8.3	< 1.11

Columns: (2) Outflow absorber velocity; (3) Outflow absorber column density; (4) ISM absorber column density

Table 6: Best-fit CO parameters

Object	Band	Continuum (FEFU)	Δv_{broad} (km s $^{-1}$)	log(N $_{CO}$) (cm $^{-2}$)	T $_{CO}$ (K)	v $_{CO}$ (km s $^{-1}$)	χ^2
V4046 Sgr	14 - 3	3.1	27 ± 1	18.55 ± 0.45	241 ± 41	-2.6	2.38 - 5.2
V4046 Sgr	14 - 4	3.5	30 ± 3	18.47 ± 0.41	295 ± 87	-6.4	5.18 - 7.55
DF Tau	14 - 3	2.8	21 ± 1	18.56 ± .38	434 ± 135	27	< 1.08
RECX 15	14 - 3	0.4	16 ± 1	18.50 ± 0.53	318 ± 131	1.0	< 1.14
RECX 15	14 - 4	0.5	20 ± 1	18.54 ± 0.57	427 ± 249	1.0	1.7 - 2.04
DM Tau	14 - 3	0.5	19 ± 4	19.21 ± 1.14	833 ± 562	35.0	< 1.18
DM Tau	14 - 4	0.8	17 ± 3	19.08 ± 0.77	775 ± 445	26.2	< 1.14
RECX 11	14 - 3	0.2	53 ± 33	18.36 ± 0.50	347 ± 158	18.0	< 1.13
HD 135344B	14 - 3	1.0	41 ± 14	18.58 ± 0.93	148 ± 66	5.0	< 1.13
DE Tau	14 - 3	0.3	26 ± 20	18.35 ± 1.03	793 ± 696	30.0	< 1.14

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Columns: (3) Local continuum for CO band; (4) FWHM of individual J emission lines; (5) CO column density; (6) CO temperature; (7) CO LSR velocity

Table 7: CO Band Summed Fluxes

Target	$C_{(14-3)}$ (FEFU)	Blue $F_{(14-3)}$ (FEFU Å) ^a	Red $F_{(14-3)}$ (FEFU Å)	$L(CO)$ (10^{28} ergs s $^{-1}$)
V4046	3.30 ± 0.68	47.36 ± 0.39	20.33 ± 0.28	33.70 ± 0.24
DFTAU	3.73 ± 1.02	4.69 ± 0.31	17.38 ± 0.40	31.25 ± 0.71
RECX15	0.54 ± 0.31	2.11 ± 0.12	3.35 ± 0.15	3.71 ± 0.13
DMTAU	2.20 ± 1.02	9.19 ± 0.41	4.21 ± 0.39	18.98 ± 0.80
RECX11	0.33 ± 0.24	2.18 ± 0.11	1.01 ± 0.10	2.17 ± 0.10
HD135344B	1.28 ± 0.73	6.38 ± 0.39	3.18 ± 0.35	13.55 ± 0.75
DETAU	1.54 ± 2.24	3.04 ± 0.52	0.99 ± 0.49	5.70 ± 1.01
GMAUR	2.82 ± 1.37	2.70 ± 0.40	0.60 ± 0.39	4.68 ± 0.79
HNTAU	2.16 ± 1.12	1.30 ± 0.29	1.43 ± 0.31	3.87 ± 0.60
DRTAU	14.79 ± 8.81	0	0	0
DKTAU	4.14 ± 4.90	4.57 ± 1.49	8.16 ± 1.74	18.02 ± 3.25
LkCa15	1.61 ± 1.50	0	0	0

^a1 FEFU Å = 10^{-15} ergs cm $^{-2}$ s $^{-1}$